

## POWERFUL SOLAR FLARES IN SEPTEMBER 2017. COMPARISON WITH THE LARGEST FLARES IN CYCLE 24

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*Solar flare activity in cycle 24 is studied. Satellite observations of x-ray fluxes from GOES-15 and UV emission lines from the SDO/EVE experiment are used. The most powerful flares of cycle 24 in classes X9.3 and X8.2 in September 2017 are compared with powerful flares in classes M5-X6.9. The times at which the fluxes in the 30.4 and 9.4 nm lines and in the 0.1-0.8 nm x-ray range begin to increase are compared for 21 of the large flares. The total energies arriving at the earth from flares in the 30.4 and 9.4 nm lines and in the 0.1-0.9 nm x-ray range,  $E_{30.4}$ ,  $E_{9.4}$ , and  $E_{0.1-0.8}$  from 25 flares during 2011 and 2012 are calculated. It is shown that the calculated energies of the flares in the analyzed lines from SDO/EVE and in the x-ray range from GOES-15 are closely interrelated.*

Keywords: sun: cycle 24: activity: flares: development of flares in lines: total flare energy

### 1. Introduction

The most powerful flares observed on the sun eject immense energy into the surrounding space, roughly a fifth of the energy released by the sun per second (for comparison, this equals all the energy released by humanity over a million years at current rates). Among the stars with activity of the solar type (e.g., UV Cet-type flare stars),

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the sun has relatively low flare activity [1,2].

The current solar activity cycle 24 is currently approaching the lowest number of solar spots and other global indices. There have been about 800 large flares in cycle 24 over the entire cycle according to an x-ray classification based on measurements by the GOES series of satellites (the classes  $>M1$  correspond to flares with amplitudes exceeding  $1 \cdot 10^{-5} \text{ W/m}^2$  in the 0.1-0.8 nm range), while there were about twice this number in the previous, more powerful cycles 22 and 23. The most powerful flare in cycle 24 up to September 2017 is considered to be in x-ray class X6.9 and took place on August 9, 2011. It does not have entirely standard characteristics for such a large flare; in particular, in terms of the amount of energy incident on the earth in the 0.1-0.8 nm x-ray range, it was not even among the first ten. The flare of August 9, 2011 was the most substantial in cycle 24 through September 2017. The features of this class X6.9 flare are discussed in detail in Refs. 3 and 4.

## 2. The flares in September 2017

On August 29, group 2673 with an area of 70 MH (1 MH is one millionth of the visible hemisphere of the sun) emerged from the eastern limb in the southern hemisphere with a total of 1 spot. After two days the area of this group fell to 60 MH, while the number of spots increased to 4 and the magnetic configuration became more complicated. On September 3, the group reached the central meridian with an area of 130 MH and 10 spots. Over the following days, the area of the group increased to 680 MH and the number of spots increased to 28, while the magnetic configuration became still more complicated. A series of solar flares began on Monday, September 4, and on that day 7 large class M flares were recorded.

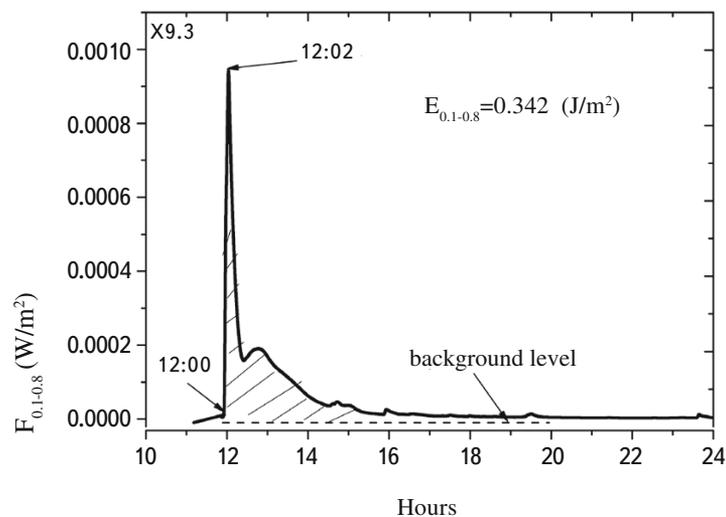


Fig. 1. The flare of Sept. 6, 2017 was the maximum in terms of x-ray classification in cycle 24, X9.3. The X axis is Greenwich time GMT.

On September 6, this group had 33 spots and an area of 880 MH, and was already in the western hemisphere near the central meridian. At 12:10 Moscow time, an X2.2 flare with a duration of 20 min was recorded in the group and at 15:02, another more powerful X9.3 flare lasting 17 min (see Figs. 1 and 2). Both flares were accompanied by proton events. Over the last twenty years, only five flares with larger amplitudes than X9.3 have been recorded, and the last of these, of class X17.0, took place almost 12 years ago (September 7, 2005).

Figure 1 is a plot of the radiative flux in the interval 0.1-0.8 nm (linear scale). The total energy arriving from the flare to the earth, calculated as the area under the shaded time dependence of the flux in the 0.1-0.8 nm range minus the background level, is indicated in the figure.

Against the background of a prolonged influx of protons from an M5.5 flare, on Sept. 6, 2017 at 12:55 GMT, protons from an X9.3 flare began to appear over a wide range of energies. For protons with energies  $\geq 100$  MeV, this was a significant event that occurred for the first time since 2014.

Figure 2 illustrates the proton event caused by the flare of Sept. 6, 2017, based on data from observations of proton fluxes with the GOES-15 satellite at energies  $\geq 10$  MeV and  $\geq 100$  MeV. It can be seen that after a maximum of protons with energies  $\geq 100$  MeV at 12:35 GMT, a gradual decrease in the proton flux began. For protons with energies  $\geq 100$  MeV, the enhanced flux of protons from this flare was superimposed almost unnoticed against the increased flux of protons from the flare of Sept. 7, after which a further increase in the flux of protons with energies  $\geq 100$  MeV set in.

On Sept. 10, 2017, at about 16 h Greenwich time, there was yet another very powerful flare of class X8.2 (see Figs. 3 and 4). It was a continuation of the very powerful X9.3 flare of Sept. 6 and was the next to most powerful since 2005, being less powerful only than its immediate predecessor of four days before. This flare, which was associated with the same active region 2673, took place essentially in the limb; on the next day, the active region 7623 went past the edge of the sun, but the hard proton source associated with this flare and located above the active

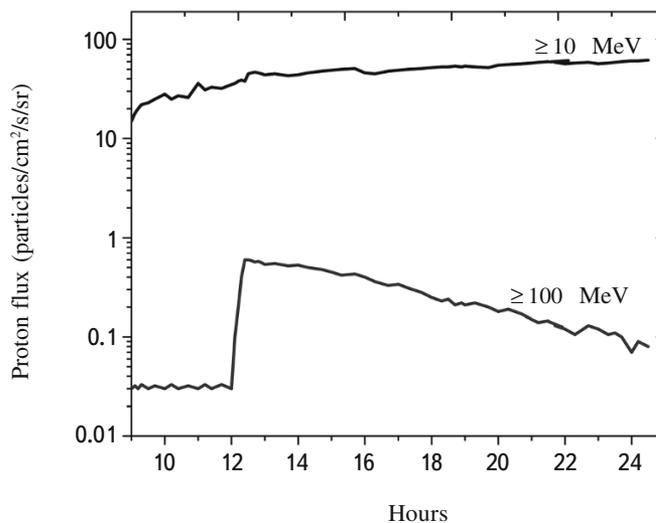


Fig. 2. The flare of Sept. 6, 2017. (top) Flux of protons with energies  $\geq 10$  MeV; (bottom) flux of protons with energies  $\geq 100$  MeV.

region in the corona, continued to shine above the limb for more than a week. The radiative flux in the 0.1-0.8 nm range is indicated in Fig. 3 (linear scale). The total energy  $E_{0.1-0.8}$  arriving on the earth from the flare is indicated in the figure; it was calculated as the area under the shaded time dependence of the flux in the 0.1-0.8 nm range.

On Sept. 10, 2017, at roughly 16:30 GMT, proton fluxes from an X8.2 flare began to appear over a wide range of energies with energies considerably higher than those produced by the X9.3 flare of Sept. 6, 2017.

The enhanced fluxes of protons in the flare of Sept. 10 reached record levels for this cycle, both for protons with energies  $\geq 10$  MeV and for the harder  $\geq 100$  MeV protons. The proton fluxes in the flare of Sept. 10 exceeded those in the flare of Sept. 6 by more than an order of magnitude (see Fig. 4) and this was a record level for all of cycle 24. The effect of these protons on the earth's atmosphere was grandiose: the magnetic storms following this flare were also of record magnitude and duration.  $k_p$ , the index characterizing the degree of perturbation of the earth's magnetic field reached 6 units over three days according to NOAA/SWPC data (see the archive at the site [http://www.n3kl.org/sun/noaa\\_archive](http://www.n3kl.org/sun/noaa_archive)).

It is likely that the effect of this proton flare on the earth's atmosphere would have been even greater if the proton source had not moved together with the flare region beyond the sun's limb.

We emphasize that these powerful flares have taken place against the background of a solar minimum during cycle 24. The active region 7623 existed for another two revolutions of the sun, but further significant flare activity in this region was not observed.

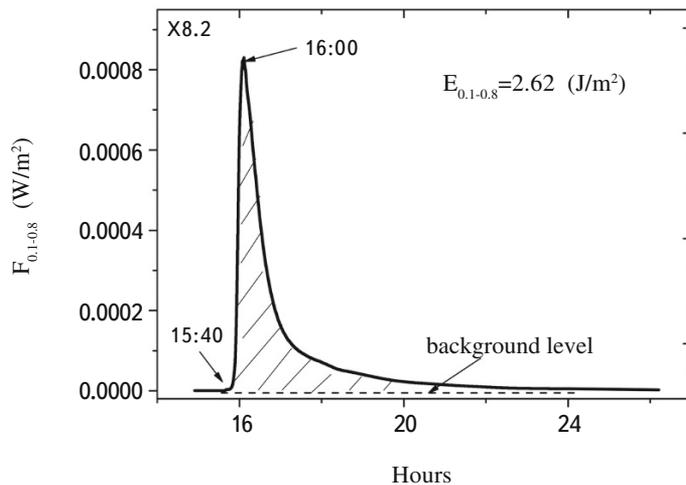


Fig. 3. The flare of Sept. 10, 2017, one of two highest in terms of x-ray classification in cycle 24. The X axis is Greenwich time GMT.

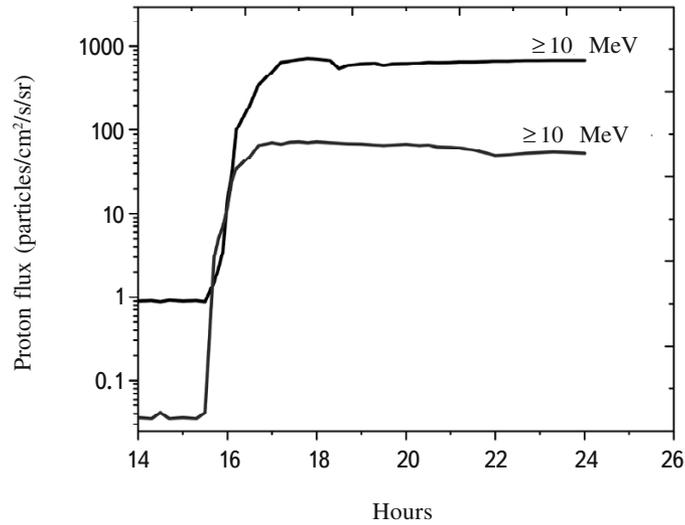


Fig. 4. The flare of Sept. 10, 2017. From the top: the flux of protons with energies  $\geq 10$  MeV; from the bottom, the flux of protons with energies  $\geq 10$  MeV.

### 3. Large flares in cycle 24. The time of onset and maximum of a flare in the 30.4 nm and 9.4 nm lines and in the interval 0.1-0.8 nm

As opposed to cycles 21-23, weak flare activity was observed in cycle 24: a total of 133 X-class  $>M5.0$  flares, of which 49 were X-class  $>X1$ . The largest number of these flares was observed near the first and second maxima of the cycle (of the 10 largest flares, only two took place on the falling branch: these are the largest flares in cycle 24 which occurred on September 6 and September 10, 2017) [5,6]. By comparison with cycles 21-23, the largest X-class x-ray flares  $>X15$  were observed on the falling branches of cycles 21 and 23, as well as in the maximum of cycle 22 [7-9]. Observations of flares on satellites with high temporal resolution can be used to study the delay in the onset time of flares in lines belonging to the lower part of the sun's atmosphere (chromosphere and transition region) and in the upper part (corona). SDO/EVE observations in extreme ultraviolet lines are available with averaging over 1 min. GOES-15 observations over 0.1-0.8 nm and 0.05-0.4 nm are available with a time resolution of 2.5 s. Thus, for each of the flares being studied, we can obtain data on the times of the flare onset and maximum in the chosen lines.

Figure 5 shows the x-ray fluxes in 0.1-0.8 nm from GEOS-15 and in 5 lines from SDO/EVE for a large class M8.8 flare that lasted more than 4 h. This is one of the largest flares in terms of total energy ( $E_{0.1-0.8} = 0.389 \text{ J/m}^2$ ) in cycle 24. This figure shows the typical times of the inflections corresponding to the onset of a flare and its maximum (two maxima can be seen in both the 30.4 and 13.3 nm lines). The times of flare onset and the maximum indicated in Fig. 5 were refined directly from the time series used to construct the plots. It is clear that the typical

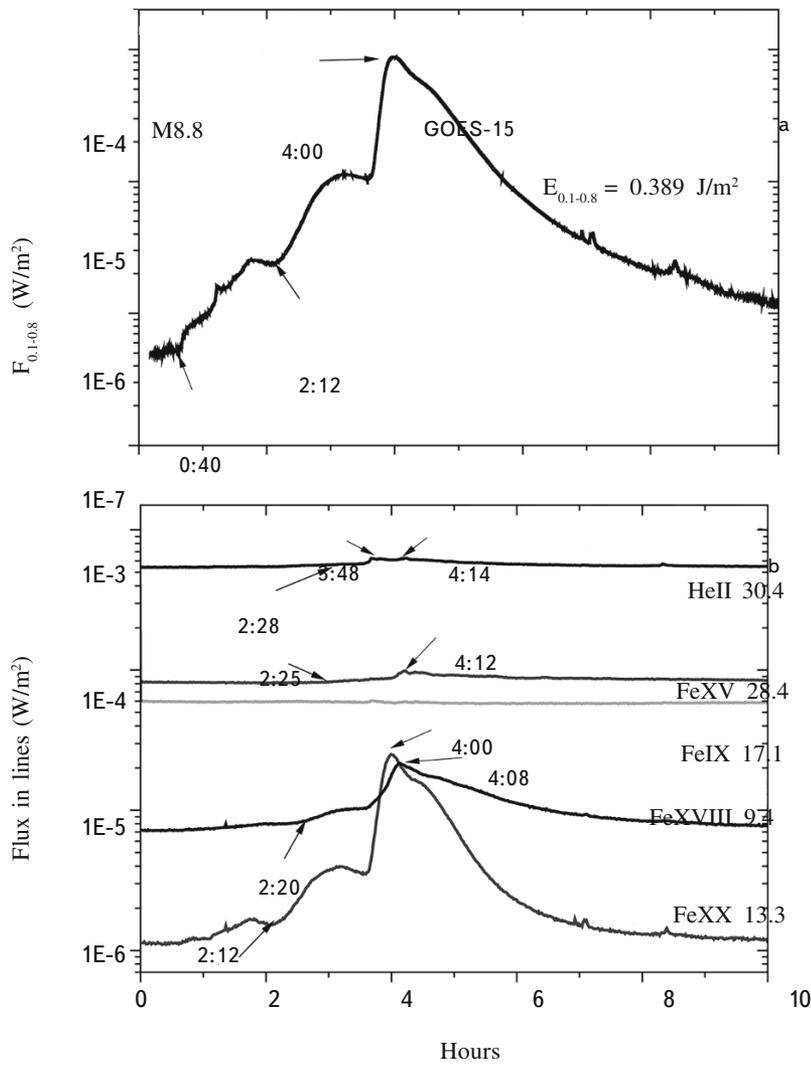


Fig. 5. The M8.8 flare of January 23, 2012. (a) GOES-15 flux in 0.1-0.8 nm; (b) SDO/EVE fluxes in UV lines.

time profile of a flare in the GOES-15 data for 0.1-0.8 nm corresponds best to the time profile of the FeXX 13.3 nm coronal line and only slightly less well to the time profile of the FeVIII 9.4 nm coronal line. The maximum concentrations of FeVIII and FeXX ions are observed in the upper corona for  $T \sim 10^7$  K and the emission over 0.1-0.8 nm is formed roughly in this part of the corona [10,11].

For the flare of January 23, 2012, shown in Fig. 5, the corona lines within the 0.1-0.8 nm range (2:12-2:20) become more intense earlier and the chromospheric lines in the transition region do so later (2:25-2:28).

#### 4. Relationship of the total energy radiated by flares in the 0.1-0.8 nm range to the primary energy release region

For further analysis of the region of primary energy release for the 21 flares we have chosen the lines at 30.4 and 9.4 nm. On one hand, these lines are formed at different altitudes in the sun's atmosphere (30.4 nm is from the

transition region and 9.4 nm is a corona line). On the other hand, these lines have additional useful properties: the 30.4 nm line is one of the strongest in the UV and plays a significant role in the formation of the earth's ionosphere, while the 9.4 nm line is very sensitive to the sun's flare activity and is an order of magnitude more intense than the 13.3 nm line. The series of observations in the 30.4 and 9.4 nm lines was broken in May 2014 because of failure of some of the measurement instrumentation on the SDO satellite.

We collected data on 21 flares of class >M5 (observations in 2011-2012) and analyzed them in the 30.4 and 9.4 nm lines an the 0.1-0.8 nm range.

14 of these 21 large flares (indicated by the letter K in the first column of Table 1) were similar to the flare of January 23, 2012, for which the intensification of the fluxes began with the 9.4 nm corona line and the 0.1-0.8

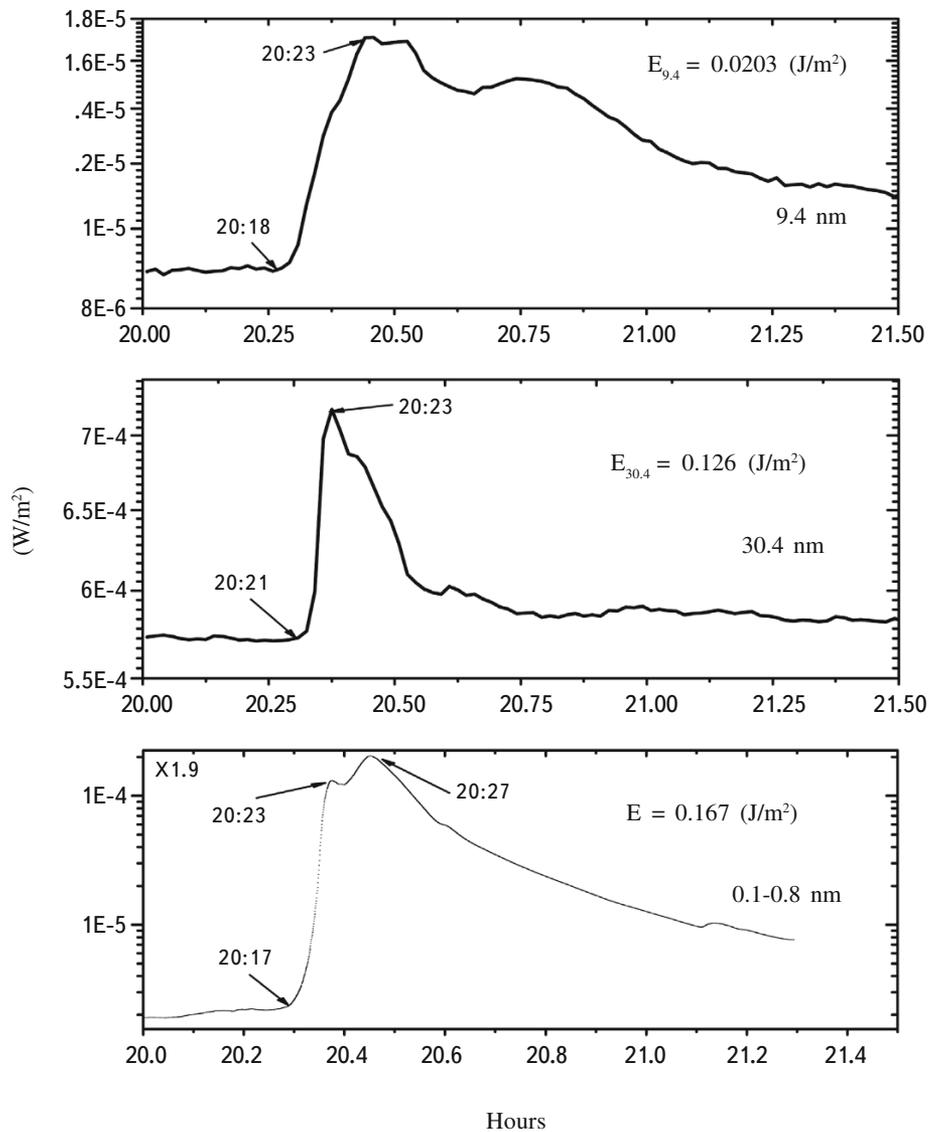


Fig. 6. The X1.9 flare of November 3, 2011.

nm range. As an example of such a flare, we consider the X1.9 flare in the corresponding UV lines and x-ray range which took place on November 3, 2011 (see Fig. 6).

Figure 6 shows the class X1.9 flare of November 3, 2011, which lasted a fairly long time (about one and a half hours) and had a high total energy incident on the earth in the 0.1-0.8 nm range,  $E_{0.1-0.8} = 0.167 \text{ J/m}^2$ . The first part of Fig. 6 shows the same flare in the 30.4 and 9.4 nm lines. A comparison of the temporal fluxes from the flare of November 3, 2011 in the 0.1-0.8 nm range and the 30.4 and 9.4 nm lines showed that the initial intensification of the flare took place in the corona, and after 3-4 min, in the transition region. Here the flare maximum coincides in time in both lines and in the 0.1-0.8 nm range.

An analysis of 6 large flares out of the 21 studied here (indicated by the symbols K for corona and TR for transition region in Table 1) showed that for these flares the intensification of the fluxes occurs roughly at the same time (to within 1 min) in the observed lines and x-ray interval. As an example of these flares, we examine the flare of March 9, 2011 (see Fig. 7). This flare is not very long (less than an hour) and, therefore, not highly energetic.

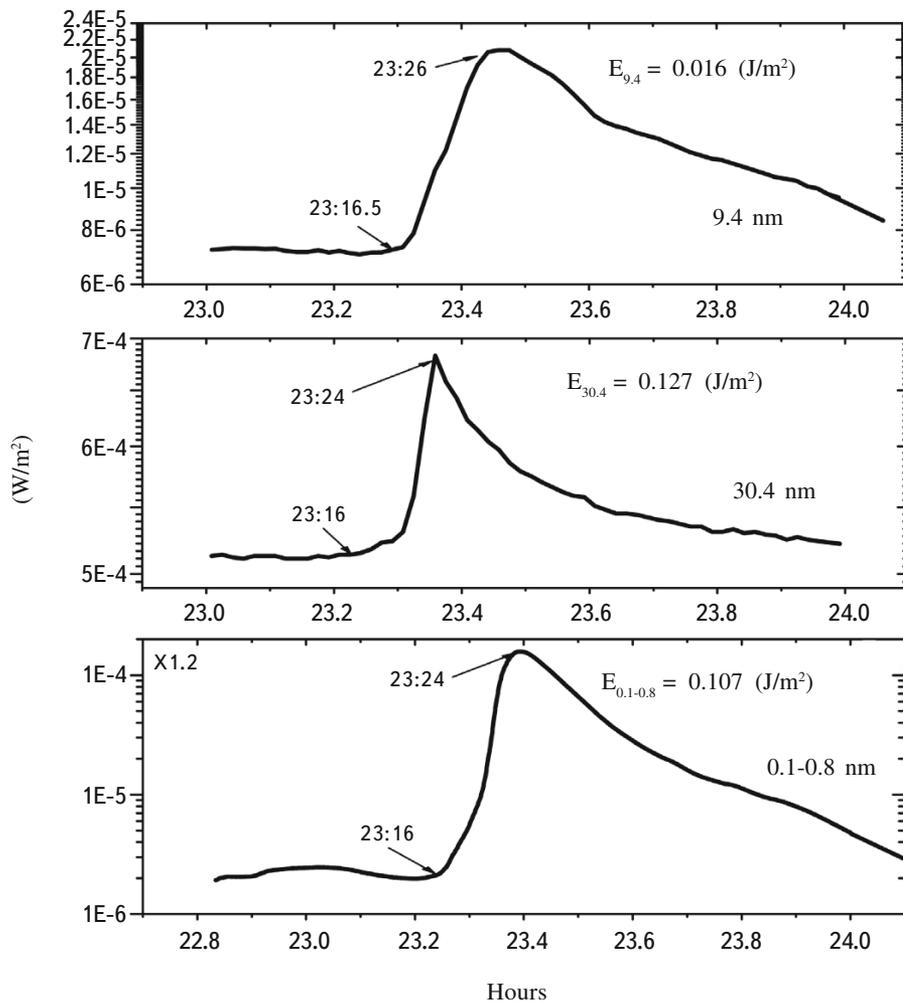


Fig. 7. The X1.2 flare of March 9, 2011.

A comparison of the time variations in the fluxes from the flare of March 3, 2011, in the 0.1-0.8 nm range and the 30.4 and 9.4 nm lines showed that the initial intensification of the flare is essentially simultaneous in the corona and in the transition region. The flare maximum is slightly delayed in the 9.4 nm line.

It can be seen clearly in one of the smallest M3.5 flares that occurred on February 24, 2011, and lasted about an hour (indicated by the symbol TR in Table 1), that the intensification begins in the lower part of the atmosphere in the 30.4 nm line in the transition region, and later, after 2-3 min, in the coronal line at 9.4 nm and in the 0.1-0.8 nm range (see Fig. 8).

Thus, a relationship shows up between the energy of a flare and the region from which it initially releases energy. It is evident that more data on flares including weaker flares beginning with class  $\geq M1$  will be needed to confirm the assumption that the initial energy release region for low energy flares is not in the corona, but in the chromosphere and transition region.

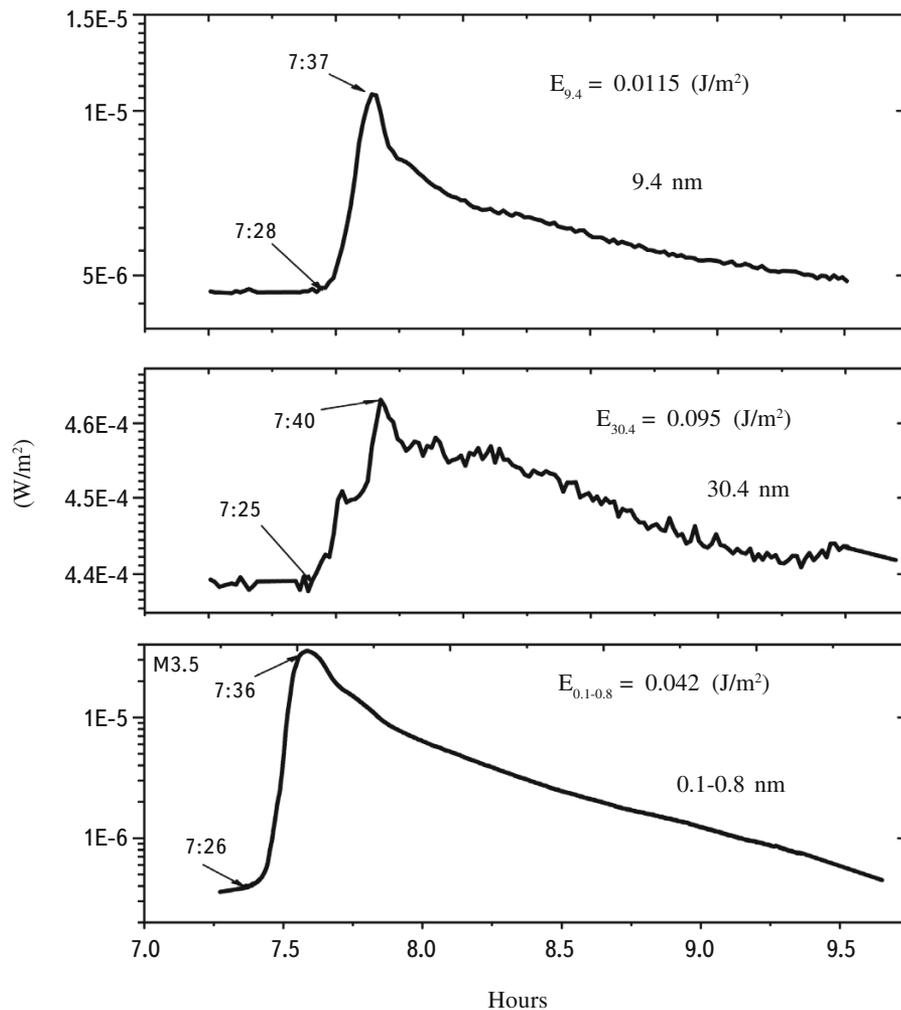


Fig. 8. The M3.5 flare of February 24, 2011.

The parameters of the 21 flares that we examine here are listed in Table 1. Data on the background flux level and on the magnitude and time of the beginning and maximum of the flares in the 0.1-0.8 nm range are listed for each flare.

The flares in our sample are among the relatively large flares during 2011-2012. We emphasize that in Table 1 they are arranged in order of decreasing total energy emitted in the 0.1-0.8 nm range. The energy  $E_{0.1-0.8}$  in that range was calculated for each flare as the area under the time dependence of the flux from the flare minus the background flux. For the flares of September 6 and September 10, 2017, these are the shaded areas under the plot of the total energy  $E_{0.1-0.8}$  shown in Figs. 1 and 3 (there were no UV data from SDO/EVE during 2017).

In Table 1, we have identified the flares in terms of the primary energy release region.

The symbol K in column 1 of Table 1 corresponds to flares where the intensification of the flare fluxes began

TABLE 1. Parameters of the 21 Large Solar Flares of 2011-2012 Based on Observations from GOES-15 in the 0.1-0.8 nm Range

Date of flare/ Region of initial energy release	Background level (W/m <sup>2</sup> )/time	Onset of flare (W/m <sup>2</sup> )/ time	Flare maximum (W/m <sup>2</sup> )	$E_{0.1-0.8}$ (J/m <sup>2</sup> )
22.09.11/K	1.05E-6/09:00	7.4E-6/10:35	1.5E-4/11:00	0.756
15.02.11/K	1.12E-6/01:44	1.25E-6/01:47	2.3E-4/01:56	0.2628
09.08.2011/K	8E-7/07:45	1E-6/08:00	7.4E-4/08:05	0.2574
08.03.2011/K, TR	1.4E-6/18:05	2.0E-6/18:10	4.45E-5/18:27	0.0806
07.03.2011/K	2.2E-6/19:20	3.73E-6/20:12	5.0E-5/20:42	0.179
03.11.2011/K	2.06E-6/20:06	2.35E-6/20:17	2.04E-4/20:27	0.168
03.08.2011/K	1.3E-6/17:45	2E-6/18:10	4.44E-5/18:28	0.148
24.09.2011/K	1.8E-6/09:31	4.00E-6/9:34	1.91E-4/09:40	0.143
06.09.2011/K	1.8E-6/22:03	2.02E-6/22:13	2.16E-4/22:20	0.118
04.08.2011/K	6E-7/03:43	2E-6/03:44	9.5E-5/03:57	0.112
07.09.2011/K, TR	6.0E-7/22:13	3.0E-6/22:35	1.8E-4/22:38	0.1008
08.03.11-1/K	1.87E-6/10:33	3E-6/10:30	5.42E-5/10:44	0.0803
09.03.11/K, TR	2.4E-6/23:00	3.0E-6/23:16	1.58E-4/23:23	0.107
13.02.2011/K	6.5E-7/17:26	2.0E-6/17:32	6.9E-5/17:38	0.072
25.09.2011/K	3E-6/14:55	6E-6/15:22	3.7E-5/15:43	0.064
08.09.2011/K	5E-7/15:31	1E-6/15:36	6.76E-5/15:46	0.059
18.02.2011/K	1.2E-6/09:44	3.0E-6/10:11	7.38E-5/10:12	0.053
26.12.2011/TR	1.13E-6/20:00	1.6E-6/20:15	2.4E-5/20:30	0.0422
24.02.2011/K, TR	4.0E-7/07:22	1.0E-6/07:26	3.55E-5/07:35	0.0421
14.03.2011/K, TR	5.0E-7/19:32	3.0E-6/19:46	4.43E-5/19:52	0.022
31.12.2011/K, TR	6.8E-7/13:09	1.7E-6/13:12	2.49E-5/13:15	0.0126

in the corona. The symbols K and TR together correspond to flares where the intensification of the flare fluxes took place roughly simultaneously in the corona and in the transition region.

The class M3.5 flare of February 24, 2011, with the TR symbol alone corresponds to an earlier intensification of the flux in the 30.4 nm line of the transition region, followed by the 9.4 nm corona line and the 0.1-0.8 nm interval.

It is clear that the energy  $E_{0.1-0.8}$  (fifth column of Table 1) from a flare is associated with the initial phase of development of the flare process, i.e., the height in the atmosphere where the intensification in the lines begins. Table 1 shows that for the more energetic flares, the intensification begins in the corona. For the less energetic ones, the flare begins in the corona or simultaneously in the corona and the transition region. For one of the weakest flares in this sample (No. 19 of the 21), the intensification of the flare flux begins in the lower part of the atmosphere. A similar dependence on the class of a flare (its peak amplitude, column 4 of Table 1) is not so evident, since powerful x-ray flares of class X can be of short duration, while longer lasting class M flares can have large values of  $E_{0.1-0.8}$ .

## 5. Interrelationship of the energies radiated by flares in the 30.4 and 9.4 nm lines and in the 0.1-0.8 nm range

Table 2 lists the total energies incident on the earth in the 30.4 and 9.4 nm lines and in the 0.1-0.8 nm range from 25 large flares. The energies  $E_{0.1-0.8}$ ,  $E_{30.4}$ , and  $E_{9.4}$  are calculated for each flare as the area under the time dependence of the flux emitted from the flare minus the background flux. Besides the date of the flares, column 1 indicates the flares which were accompanied by ejection of protons (pr).

Table 2 shows that the energy of a flare is related to whether the flare is accompanied by a proton event.  $E_{0.1-0.8}$  approximately equal to  $0.1 \text{ J/m}^2$  is a typical value below which flares are not associated with subsequent ejection of protons.

Figure 9 shows the relationship among the energies  $E_{0.1-0.8}$ ,  $E_{30.4}$ , and  $E_{9.4}$ . We see a fairly close relationship between the calculated energies. Evidently, the relationship shown in Fig. 9 can be refined by adding new flares with large energies (the flares that took place up to mid 2014, when complete observations of the UV lines were made with SDO/EVE at the same time as the GOES-15 observations). Since these GOES 0.1-0.8 nm observations are essentially available in real time, while the SDO/EVE UV observations were cut off in mid 2014, the relationship between the energies is important for estimating the energies emitted in the lines based on  $E_{0.1-0.8}$ .

## 6. Conclusions

1. Cycle 24 is characterized by low flare activity. The number of large flares was roughly half that in cycles 22 and 23. The largest flares in September 2017 were of classes X9.4 and X8.3, while class X13-X17 flares occurred in cycle 23.

2. The most important characteristic of the flares is the total energy from them incident on the earth's surface per  $\text{m}^2$  and is the input parameter for analyzing the effect of a flare on the earth's upper atmosphere.

TABLE 2. Calculated Energies for 25 Large Flares During 2011-2012  
Based on Data from GOES-15 and SDO/EVE

Date of flare	$E_{0.1-0.8}$ (J/m <sup>2</sup> )	$E_{30.4}$ (J/m <sup>2</sup> )	$E_{9.4}$ (J/m <sup>2</sup> )
12.07.2012 pr	0.792	0.956	0.118
22.09.2011 pr	0.756	0.66	0.097
23.01.2012 pr	0.389	0.91	0.1123
15.02.2011 pr	0.263	0.27	0.028
09.08.2011 pr	0.257	0.099	0.0187
07.03.2011 pr	0.179	0.274	0.024
03.11.2011 pr	0.168	0.126	0.0203
03.08.2011 pr	0.148	0.183	0.0205
24.09.2011 pr	0.143	0.11	0.021
31.12.2011 pr	0.126	0.035	0.0031
23.10.2012 pr	0.1188	0.084	0.0195
06.09.2011 pr	0.118	0.124	0.0116
04.08.2011 pr	0.112	0.237	0.256
09.03.2011 pr	0.107	0.127	0.016
07.09.2011 pr	0.101	0.197	0.0154
08.03.2011-2	0.081	0.012	0.0179
08.03.2011-1	0.080	0.081	0.0124
13.02.2011	0.072	0.101	0.082
25.09.2011	0.064	0.0761	0.0078
08.09.2011 pr	0.059	0.153	0.009
18.02.2011	0.053	0.103	0.006
26.12.2011	0.042	0.065	0.0054
24.02.2011	0.042	0.095	0.0115
14.03.2011	0.022	0.057	0.0045
20.10.2012	0.064	0.594	0.0098

The flare energies in the 30.4 nm transition region line and the 9.4 nm corona region line are closely related to one another and to the x-ray flux in the 0.1-0.8 nm region, while the amplitudes of the flares in these lines and the x-ray range are not distinctly related. The energy of a flare determines whether it is a proton flare: beginning with  $E_{0.1-0.8} \geq 0.1$  J/m<sup>2</sup>, the flares are accompanied by proton emission.

3. The energy of a flare is related to the region in which initial intensification of the flux of flare emission takes place (based on an analysis of the 30.4 nm transition region line and the 9.4 nm corona line, as well as the emission in the 0.1-0.8 nm range): when the flare energy is higher, it is more probable that the region of the initial intensification is in the corona.

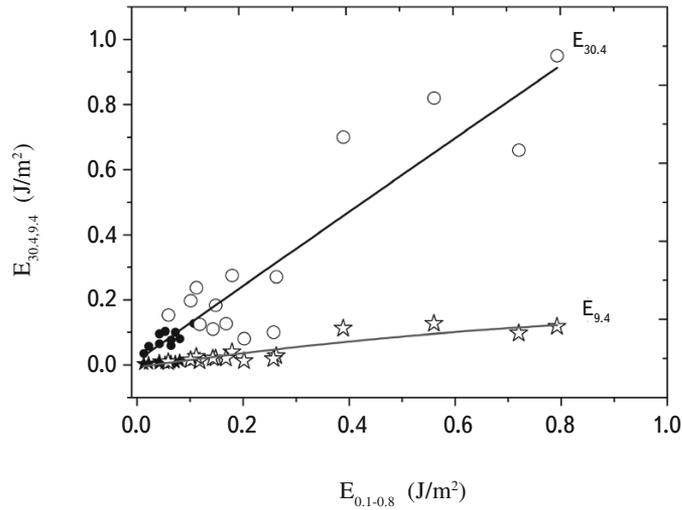


Fig. 9. The relationship between the total energy radiated by a flare in the UV lines at 30.4 and 9.4 nm and in the 0.1-0.8 nm x-ray range. The circles correspond to the 30.4 nm line and the stars to the 9.4 nm line. The large hollow circles and stars indicate that the flares were accompanied by proton ejection.

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